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Bias and Conditional Mass Function of Nonspherical Dark Halos

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ABSTRACT

Nonspherical collapse is modelled, under the Zeldovich approximation, by six-dimensional random walks of the initial deformation tensor field. The collapse boundary adopted here is a slightly-modified version of that proposed by Chiueh and Lee (2001). Not only the mass function agrees with the fitting formula of Sheth and Tormen (1999), but the bias function and conditional mass function constructed by this model are also found to agree reasonably well with the simulation results of Jing (1998) and Somerville et al. (2000), respectively. In particular, by introducing a small mass gap, we find a fitting formula for the conditional mass function, which works well even at small time intervals between parent and progenitor halos during the merging history.

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1. Introduction

In a hierarchical model, the gravitationally bound objects of the universe formed from mergers of smaller objects into larger ones. An issue of fundamental importance is to understand how gravitational structures, such as galaxies and clusters, are distributed as a function of mass and evolving time. Press & Schechter (1974) provided a simple analytical model (hereafter PS model) to describe the mass distribution of dark halos (mass function). This theory is based on the assumption that the process of gravitational collapse can be approximated by spherical symmetry and it occurs when the mean density contrast of pre-collapse regions predicted by the linear theory reaches a threshold value δ_c . Bond et al. (1991) introduced the excursion set formalism and extended the PS theory to calculations of the conditional mass function, which is closely related to the issue of the linear bias and halo merging histories. Mo & White (1996) used the extended Press-Schechter theory (hereafter EPS) to address the large-scale bias, which concerns how the collapsed halos trace the underlying dark matter. Moreover, the EPS theory is also useful for construction the merger history of dark halos (Bower 1991; Lacey & Cole 1993, 1994 hereafter LC93, LC94), which concerns the probability that a halo of mass M_2 at time t_2 has a progenitor halo of mass M_1 at earlier time t_1 .

The PS or EPS theory has been tested against N-body simulations (Efsthathiou et al. 1988; LC94), and the comparisons suggest that the PS model overestimates the number of halos in low-mass regimes and underestimates it in high-mass regimes. This discrepancy indicates that the spherical collapse model used by the PS theory may be too simplistic. Many researchers have taken the nonspherical collapse process into account to correct this problem (e.g., Bond & Myers 1996; Monaco 1995; Lee & Shandarin 1998; Chiueh & Lee 2001; Sheth, Mo and Tormen 2001). Most of these works are still based on the main idea of the PS theory, and further incorporate either the dynamics of ellipsoids or the Zeldovich approximation (Zeldovich 1970).

The main aims of this paper are to investigate the large-scale bias and conditional mass function by using the nonspherical collapse boundary (hereafter NCB) proposed by Chiueh and Lee (2001). We also test whether the results derived from nonspherical collapse boundary can provide better statistical predictions of dark halos than the EPS model does.

The outline of the paper is as follows. In § 2, we summarize the PS and EPS theory, and briefly review the Zeldovich approximation and the nonspherical collapse boundary proposed by Chiueh & Lee (2001). In § 3, we show the algorithm and the results of mass function and large-scale bias of our model. In § 4, the conditional mass function is derived and the comparison with N-body simulation is presented. In § 5, the conclusion is given.

2. (Extended) Press-Schechter Formalism

2.1. Mass function

The Press-Schechter theory (1974) assumes that the initial density fluctuations are Gaussian random fields. Let $\delta_R(\vec{x})$ denotes the overdensity field of at position \vec{x} in the comoving coordinate when the density distribution is smoothed over scale R by a window function W_R . The relation is

$$\delta_R(\vec{x}) = \int \delta(\vec{x}') W_R(\vec{x} - \vec{x}') d^3 \vec{x}', \quad (1)$$

with a squared variance

$$S = \sigma^2(M) = \langle |\delta_R(\vec{x})|^2 \rangle = \frac{1}{(2\pi)^3} \int p(k) \widetilde{W}_R^2(k) 4\pi k^2 dk, \quad (2)$$

where $p(k)$ is the power spectrum of density fluctuations and $\widetilde{W}_R(k)$ the Fourier transform of the window function $W_R(\vec{x} - \vec{x}')$. The squared variance S increases as mass M decreases, and in this paper, we will often use S to represent the mass scale.

Bond et al. (1991) extends the original PS method and uses random walks (or called excursion set) to address the mass function problem. When variance S equals zero, the density contrast δ is set to zero. As S increases, δ changes randomly. When δ first hits the threshold value δ_c , the random walk stops and the region of mass scale S , over which δ is smoothed, collapses into a bound object. This is equivalent to a diffusion process with an absorbing boundary at $\delta = \delta_c$, which has been solved by Chandrasekhar (1943). Thus, one can calculate the probability of a trajectory having its first hit at the threshold δ_c in the mass range between S and $S + dS$ to represent the mass function, which is described by

$$f(S, \delta_c) dS = \frac{\delta_c}{\sqrt{2\pi} S^{3/2}} \exp\left(-\frac{\delta_c^2}{2S}\right) dS, \quad (3)$$

where the functional form $f(x)$ represents the differential probability in the range between x and $x + dx$. Moreover, the (comoving) number density of halos with mass M is

$$\frac{dn}{dM} dM = \frac{\rho_b}{M} f(S, \delta_c) \left| \frac{dS}{dM} \right| dM, \quad (4)$$

where ρ_b is the mean mass density of the universe.

For simplicity, we define a variable $\nu \equiv \delta_c/S^{1/2}$ and one can easily show that the mass function can be represented as

$$\nu f(\nu) d\nu = \sqrt{\frac{2}{\pi}} \nu \exp\left(-\frac{\nu^2}{2}\right) d\nu. \quad (5)$$

Here δ_c is often chosen as 1.68647, which is predicted by the spherical collapse model. The main advantage of the excursion set approach is that while the mass function depends on the variance S and the value of collapse threshold δ_c , the cosmological models enter this problem only through the relation between S and the mass M .

Upon comparing with N-body simulations, one finds that the EPS formalism overestimates low-mass and underestimates high-mass halos. Of course one may vary the value of the collapse threshold δ_c , but it is not possible to match the simulation's result well for all mass ranges.

2.2. Conditional mass function

In the picture of the excursion set theory, the conditional probability represents the fraction of trajectories first hitting a higher threshold δ_{c1} among those trajectories that start from a lower threshold δ_{c2} . In the framework of EPS theory, Lacey & Cole (1993) proposed that the conditional probability can assume a modified form where the origin of the mass function $(0, 0)$ shifts to (S_2, δ_{c2}) . Therefore, the conditional probability is simply

$$f(S_1, \delta_{c1}|S_2, \delta_{c2})dS_1 = \frac{(\delta_{c1} - \delta_{c2})}{\sqrt{2\pi}(S_1 - S_2)^{3/2}} \exp\left[-\frac{(\delta_{c1} - \delta_{c2})^2}{2(S_1 - S_2)}\right]dS_1. \quad (6)$$

Here, it demands $S_1 > S_2$ and $\delta_{c1} > \delta_{c2}$.

Conditional probability is useful for investigation of large-scale bias as well as merger events of halos. In the first case, the physical meaning of δ_{c2} is not the collapse threshold at some specific time, but denotes the mean density contrast of a large uncollapsed region of mass scale S_2 . Thus we shall change δ_{c2} into δ_2 to avoid confusion. If we define the bias parameter b to be the fraction of excess dark halos with mass scale S_1 that can be found in a larger mass scale S_2 , then under the limit $S_2 \gg S_1$ and $\delta_2 \ll \delta_{c1}$, Mo & White (1996) use the EPS formalism to derive that

$$b = \frac{(\delta_{c1}/S_1^{1/2})^2 - 1}{\delta_{c1}} = \frac{\nu_1^2 - 1}{\delta_{c1}}, \quad (7)$$

where $\nu_1 \equiv \delta_{c1}/S_1^{1/2}$. For convenience, we define a characteristic mass M^* , which satisfies the requirement $S(M^*) = \delta_{c1}^2$. When $S_1 < \delta_{c1}^2$, i.e., $M_1 > M^*$, the bias factor b is positive, which means the density of halos with masses greater than M^* is larger in the denser region of dark matter. On the other hand, when $S_1 > \delta_{c1}^2$, or $M_1 < M^*$, the opposite holds.

In the case of merger events, δ_{c2} denotes the collapse threshold at a later time t_2 (as compared with t_1). Again, the meanings of notations change: $\delta_{c1} \rightarrow \delta_c(z_1)$ and $\delta_{c2} \rightarrow \delta_c(z_2)$,

the collapsed thresholds at redshifts z_1 and z_2 , respectively. Thus, $f(M_1, \delta_{c1}|M_2, \delta_{c2})dM_1$ represents the mass fraction of M_2 at redshift z_2 that is in the form of collapsed halos of mass M_1 in an earlier epoch z_1 .

2.3. Zeldovich approximation

Zeldovich (1970) provided a good way to describe the development of density perturbations into the nonlinear regime. Let $a(t)$ denotes the uniform expansion of the background. The physical coordinate \vec{r} and the comoving coordinate \vec{x} are simply connected by

$$\vec{r} = a(t)\vec{x}. \quad (8)$$

In the Zeldovich approximation, the motion of a particle is described by

$$\vec{x} = \vec{q} + b(t)\vec{p}(\vec{q}), \quad (9)$$

where \vec{q} is the Lagrangian coordinate (initial location) of a particle. The second term is the perturbation of a particle in the Lagrangian coordinate. From the conservation of mass condition, we have

$$\rho(\vec{x}, t)d^3\vec{x} = \bar{\rho}(\vec{q}, t)d^3\vec{q} = \bar{\rho}(\vec{q}, 0)d^3\vec{q}. \quad (10)$$

It implies

$$\rho(\vec{x}, t) = \bar{\rho} \frac{d^3\vec{q}}{d^3\vec{x}} = \bar{\rho} \frac{1}{\det(\frac{d^3\vec{x}}{d^3\vec{q}})} \equiv \bar{\rho} \frac{1}{\det \hat{J}}. \quad (11)$$

The tensor \hat{J} in the denominator is the sum of an identity tensor and a deformation tensor \hat{D} . Let $p(\vec{q}) = -\nabla_{\vec{q}}\Phi(\vec{q})$, then

$$D_{ij} = b(t) \frac{\partial p_i}{\partial q_j} = -b(t) \frac{\partial^2}{\partial q_i \partial q_j} \Phi. \quad (12)$$

Because \hat{D} is real and symmetric, it has three real eigenvalues, $\Lambda = (-\lambda_1, -\lambda_2, -\lambda_3)$, at every spatial location. Thus, Eq. (11) can be rewritten as

$$\rho(\vec{x}, t) = \bar{\rho} \frac{1}{(1 - \lambda_1)(1 - \lambda_2)(1 - \lambda_3)}. \quad (13)$$

When λ 's are small, the density contrast δ is

$$\delta = \frac{\rho(\vec{x}, t) - \bar{\rho}}{\bar{\rho}} \sim (\lambda_1 + \lambda_2 + \lambda_3). \quad (14)$$

The sign of the value of λ determines the divergence (negative λ) or convergence (positive λ) of the local motion. Some examples illustrate the physical meaning of Eq. (13). If $\lambda_1 = \lambda_2 = \lambda_3$, when all λ 's approach unity the density becomes infinity, and it implies the spherical collapse case. If a certain λ_i grows to unity before the other two λ 's do, then the local ellipsoid will collapse into a pancake. It is easy to realize that the latter be much more probable than the former, and therefore the gravitational collapse should be nonspherical.

Since the deformation tensor \hat{D} is symmetric, it has 6 degrees of freedom: d_{11} , d_{22} , d_{33} , d_{12} , d_{23} and d_{13} . The first three represent the diagonal terms, and the rest represent the off-diagonal ones. In addition, one can construct 6 independent Gaussian variables, namely, $Y = (y_1, y_2, y_3, y_4, y_5, y_6)$. The 6 components of the deformation tensor can be interpreted as linear combinations of $y_1 \sim y_6$:

$$d_{11} = -(y_1 + \frac{3}{\sqrt{15}}y_2 + \frac{1}{\sqrt{5}}y_3)/3, \quad (15a)$$

$$d_{22} = -(y_1 - \frac{2}{\sqrt{5}}y_3)/3, \quad (15b)$$

$$d_{33} = -(y_1 - \frac{3}{\sqrt{15}}y_2 + \frac{1}{\sqrt{5}}y_3)/3, \quad (15c)$$

$$d_{12} = \frac{1}{15}y_4, \quad (15d)$$

$$d_{23} = \frac{1}{15}y_5, \quad (15e)$$

$$d_{13} = \frac{1}{15}y_6. \quad (15f)$$

2.4. Nonspherical collapse boundary (NCB)

In order to resolve the discrepancy of the mass functions between the PS prediction and the N-body simulation, Chiueh & Lee (2001) investigate how the nonspherical collapse process may be relevant. In addition to adopting the fundamental concepts of the PS hypothesis and the excursion set approach, they have noticed that the collapse condition should contain full information of the three eigenvalues $(\lambda_1, \lambda_2, \lambda_3)$, and that the nonspherical effect is more dominant in the low-mass regime than in the high-mass regime. Taking these properties into account, they propose a nonspherical collapse boundary (NCB hereafter) having the following form:

$$\frac{\delta}{\delta_c} = (1 + \frac{r^4}{\beta})^\beta, \quad (16)$$

where β is a fitting parameter taken to be 0.15, δ_c remains the collapse threshold predicted by the spherical model, which is taken to be 1.5 by Chiueh & Lee (2001) rather than 1.69 due to the correction given by Shapiro et al. (1999), and the definition of r^2 is

$$r^2 \equiv \frac{1}{3}[(\lambda_1 - \lambda_2)^2 + (\lambda_2 - \lambda_3)^2 + (\lambda_3 - \lambda_1)^2]. \quad (17)$$

In the case of spherical collapse, the three eigenvalues equal to each other. Therefore, the quantity r can be regarded as an indicator for the degree of nonsphericity.

The form of NCB has several advantages. First, this boundary is smooth and rotational invariant in the eigenvalue $(\lambda_1, \lambda_2, \lambda_3)$ space. Second, when r is close to zero, the NCB is reduced to the spherical collapse condition as expected. Third, the variable r^2 is related to the angular momentum (Catelan & Theuns 1996), so that the NCB provides additional information about the nonspherical collapse halos.

3. Mass Function and Bias for NCB

3.1. Mass function

Here we slightly improve the nonspherical collapse boundary of Chiueh and Lee (2001) into this form:

$$\frac{\delta}{\delta_c(z)} = \left(1 + \frac{1}{\beta_1} \left(\frac{r}{\delta_c(z)/\delta_c(z_0)}\right)^4\right)^{\beta_2}, \quad (18)$$

where $\delta_c(z_0) = 1.5654$, $\beta_1 = 0.26$ and $\beta_2 = 0.159$. $\delta_c(z_0)$, β_1 and β_2 are all fitting parameters to obtain a mass function closest to the result of N-body simulation (Sheth & Tormen 1999, hereafter ST99). The eigenvalues λ satisfy the eigenvalue equation

$$\lambda^3 + (d_{11} + d_{22} + d_{33})\lambda^2 + (d_{11}d_{22} + d_{22}d_{33} + d_{33}d_{11} - d_{12}^2 - d_{23}^2 - d_{13}^2)\lambda - (d_{11}d_{23}^2 + d_{22}d_{13}^2 + d_{33}d_{12}^2 - d_{11}d_{22}d_{33} - 2d_{12}d_{23}d_{13}) = 0. \quad (19)$$

Therefore,

$$\lambda_1 + \lambda_2 + \lambda_3 = -(d_{11} + d_{22} + d_{33}), \quad (20a)$$

$$\lambda_1\lambda_2 + \lambda_2\lambda_3 + \lambda_3\lambda_1 = (d_{11}d_{22} + d_{22}d_{33} + d_{33}d_{11} - d_{12}^2 - d_{23}^2 - d_{13}^2), \quad (20b)$$

$$\lambda_1\lambda_2\lambda_3 = (d_{11}d_{23}^2 + d_{22}d_{13}^2 + d_{33}d_{12}^2 - d_{11}d_{22}d_{33} - 2d_{12}d_{23}d_{13}). \quad (20c)$$

Using Eqs. (15), (20a) and (20b), it is easy to convert r^2 into a function of y_1, y_2, y_3, y_4, y_5 , and y_6 . The result is

$$r^2 = \frac{2}{15}(y_2^2 + y_3^2 + y_4^2 + y_5^2 + y_6^2). \quad (21)$$

This neat equation has been pointed out by Sheth & Tormen (2001). Combined with $\delta = y_1$ (from Eqs. (14), (15) and (20a)), we can simulate the random process in the 6-dimensional Y -space with an absorbing boundary, instead of conducting the equivalent, but more complicated, random walks in $(\lambda_1, \lambda_2, \lambda_3)$ -space (Chiueh & Lee, 2001).

In Fig. 1 we compare the mass function derived from the Y -space diffusion with two other types of mass function. The dashed line shows the EPS’s analytical prediction, the solid line shows the Sheth & Tormen’s mass function (ST99), which describes the mass distribution in N-body simulations, and the filled circles are the data of our six-dimensional random walks. We conclude that our results agree quite well with that of N-body simulation given by ST99, except at very high mass regime. There are two reasons for us to retain this discrepancy at high mass. First, the statistics of very high mass halos is not well-controlled in N-body simulation. Therefore one can not reliably specify the number of such halos. Second, it is believed that the very high mass halos should follow spherical collapse process which is closer to the EPS method. Our results show this tendency.

3.2. Bias

The peak biasing was originally introduced by Kaiser (1984) to explain the clustering of Abell clusters, and the same idea has also been used to discuss whether the galaxy distribution traces the mass distribution (Bardeen et al. 1986). Despite focusing on specific bound objects, the bias parameter also reveals the spatial clustering of dark matter halos. One common definition of the bias is

$$\frac{\delta n}{n} = b \frac{\delta \rho}{\rho}, \quad (22)$$

or

$$b(M) = \frac{n(M, \delta_c | \delta = \delta') - n(M, \delta_c | \delta = 0)}{n(M, \delta_c | \delta = 0) \times \delta'}, \quad (23)$$

where $n(M, \delta_c | \delta)$ represents the number of bound objects with mass M which can be found in the large-scale uncollapsed region of density contrast δ . Mo & White (1996, hereafter MW96) has derived the bias factor under the limit $\delta \ll \delta_c$, as shown in Eq. (7). Another equivalent expression for the bias is given by MW96:

$$\xi_{hh}(r, M) = b^2(M) \xi_{mm}(r) \quad (24)$$

evaluated at a sufficiently large separation r , where $\xi(r)$ is the two-point correlation function and the subscripts hh , mm denote the halo-halo and matter-matter components. This definition is often used to find the bias factor in N-body simulation.

In the EPS picture, one can calculate b as:

$$b(M) = \frac{n(M, \delta_c - \delta' | \delta = 0) - n(M, \delta_c | \delta = 0)}{n(M, \delta_c | \delta = 0) \times \delta'} = \frac{1}{\sigma(M) \times n(\nu)} \frac{dn(\nu)}{d\nu} = \frac{1}{\delta_c} \frac{d(\ln n(\nu))}{d(\ln \nu)}, \quad (25)$$

which indicates that the bias function is directly related to the derivative of mass function in the log-space. However, this simplicity is no longer valid in the NCB approach due to the additional r -distribution. The bias function can be another test for the goodness of the NCB. The natural choice of the initial condition for six-dimensional random walks is that $y_1 = \delta'$, whereas y_2 to y_6 are Gaussian distributed random numbers with the same variance equal to δ' .

In Fig. 2, we show the bias functions derived from our approach for three different initial power spectra of power indices, $n = 0, -1$, and -2 . The filled circles show the results of our algorithm of Y -space random walks. The solid curves denote the fitting formula of N-body simulation results found by Jing (1998):

$$b(M) = \left(\frac{0.5}{\nu^4} + 1\right)^{(0.06 - 0.02n)} \left(1 + \frac{\nu^2 - 1}{\delta_c}\right). \quad (26)$$

As a comparison, Fig. 2 also shows the predictions of the EPS model (Mo & White 1996) and the moving-barrier model (Sheth, Mo & Tormen 2001). Although there exist statistical fluctuations in our data, the filled circles show that the bias function generated from the nonspherical collapse boundary is quite reasonable and has a similar tendency as the results derived from the moving barrier model in the low-mass scale. But the curves of our results turn to coincide with the EPS prediction of high mass. This tendency is also seen in Jing's simulations. We note that Jing's bias parameter weakly depends on the spectral index n , a property that deviates from the self-similar dynamics. If the gravitational dynamics of cold dark matter is indeed self-similar, as has long been believed to be, the most relevant N-body simulation result is that of the $n = -2$ case, due to its higher numerical accuracy at small scale. Our data agree rather well with the $n = -2$ N-body result.

4. Conditional Mass Function

According to the hierarchical scenario, dark matter halos are formed through merging with comparable-size halos and accretion of small halos. All kinds of interactions between halos result in changes of gravitational potential and consequently the behavior of gas components. Therefore, merger/accretion rates are believed to play a crucial role in the galaxy formation. Observations show that merger rates increase with redshift (Carlberg et al. 1994; Yee & Ellingson 1995), despite the different definitions of the merger rate. On the other

hand, the dependence of merger rates on redshifts and on environments have also been studied with N-body simulations (e.g., Gottlöber et al. 2001), and the results roughly agree with the observations. It is expected that the semi-analytical approach (e.g., EPS model and NCB model) can offer succinct description for merger histories. However, we must keep in mind that the semi-analytical methods deal with only the isolated halos, discarding the information of substructures within each halo. The attempt of explaining galaxy formation inside galaxy groups or clusters with semi-analytical approaches must be modified. Despite such a limitation for the cloud-in-cloud problem, the merger-tree technique is often needed to establish merger histories in the field (Kauffmann & White 1993; Somerville & Kolatt 1999). The merger-tree method is a Monte-Carlo method based on the conditional mass function. In the literature, EPS’s conditional mass function is often used to construct merger trees. However, the results indicate that there still exhibit discrepancies with the merger history extracted from N-body simulations, which is thought to be caused mainly by the EPS model rather than the specific features of merger trees (Somerville et al. 2000). Therefore, in this section, we use the NCB model to generate a new conditional mass function with the six-dimensional random walk procedure and derive a fitting formula for it.

4.1. Construction of conditional mass function

There are two scenarios, the active picture and the passive picture, to represent the diffusion process for epochs earlier than the present. The former is with a fixed collapse boundary and let the variance of some particular mass to grow with time. The latter is just the opposite which fixes the variance of mass M and lowers the collapse boundary with increasing time. In order to simulate the merging problem, we adopt the second picture for simplicity. For a given redshift z , the collapse boundary condition is given in Eq. (18). As shown in the Fig. 3, boundary 1 corresponds to the collapse condition at a time earlier than boundary 2. Note that boundary 1 should be stretched both in r and δ axes relative to boundary 2 to preserve the self-similar evolution of fluctuations. After collecting the trajectories that start from the origin and collapse at boundary 2, we continue the random walks until they hit boundary 1 and are absorbed.

Let $f(S_1, \delta_c(z_1)|S_2, \delta_c(z_2))$ denotes the fraction of trajectories hitting boundary 1 at mass scale S_1 to the total trajectories hitting boundary 2 at mass scale S_2 . The physical meaning of $f(S_1, \delta_c(z_1)|S_2, \delta_c(z_2))$ is the probability that a point in a halo with mass scale S_2 at redshift z_2 was in a halo with mass scale S_1 at redshift z_1 . As mentioned earlier, although S is the variance of the density fluctuation corresponding to a specific mass range, we use it to interpret the mass scale over which the density contrast δ is smoothed. Apparently,

S_1 is always larger than S_2 to ensure that the mass of halo increase with time. Moreover, due to the self-similar scaling property of the nonspherical collapse boundary, f does not depend on time explicitly, but only on the ratio of the separation, namely $\delta_c(z_1)/\delta_c(z_2) \equiv \kappa$. The actual time evolution of the conditional mass function is hidden in the variable μ or μ' defined below, which is a combination of mass and time. Therefore, by simulating random walks with a particular chosen $z_2 = 0$, we are able to construct the conditional mass function which can also be applied to any other z_2 .

Recalling the Pess-Schechter's boundary, its conditional probability is given in Eq. (6), or in a concise form,

$$\mu f(\mu) d\mu = 2 \times \left(\frac{\mu^2}{2\pi}\right)^{1/2} \exp\left(-\frac{\mu^2}{2}\right) d\mu, \quad (27)$$

where $\mu \equiv (\delta_c(z_1) - \delta_c(z_2))/(S_1 - S_2)^{1/2} = \Delta\delta_c/(\Delta S)^{1/2}$. Obviously, for given z_1 and z_2 , the conditional probability is a function of only the difference of the variance S in EPS picture. However, the conditional probability actually depends also on the initial mass scale S_2 in our scenario. This is due to that the separation of two boundaries is S_2 -dependent, as is obviously indicated in Fig. 3.

Here we simulate several conditions for different boundary separations: $\kappa = 1.1, 1.3, 2.0, 3.0$ and 4.0 with $z_2 = 0$. It helps the construction of a new conditional mass function by considering the following two extreme cases: When κ is small (the two boundaries are close), the conditional mass function is expected to be close to the extend Press-Schechter's case because the two boundaries are almost parallel and most points collapsing at the first boundary will soon collapse in few steps at the second boundary. On the other hand, when κ is very large, points in the second boundary relative to the first boundary are like that they are located at the origin, so that the conditional mass function should behave like the Sheth and Tormen's mass function (hereafter ST mass function). We find that our simulation data can be fitted by the following function:

$$\mu' f(\mu') d\mu' = 2A(\kappa) \left(1 + \frac{1}{\mu'^{2q}}\right) \left(\frac{\mu'^2}{2\pi}\right)^{1/2} \exp\left(-\frac{\mu'^2}{2}\right) d\mu', \quad (28)$$

where

$$\mu' \equiv \frac{[\delta_c(z_1) - \delta_c(z_2)]\varepsilon(S_2, \kappa)}{(S_1 - S_2)^{1/2}}, \quad (29)$$

$$A = 0.322 + \frac{0.178}{\kappa}, \quad (30)$$

$$\varepsilon(S_2, \kappa) = \varepsilon(x) = 0.036x^4 - 0.309x^3 + 0.944x^2 - 1.060x + 1 \quad (31)$$

with $x \equiv (\sqrt{S_2}/\delta_c(z_2) - 0.25)/\kappa$, and $q(\kappa)$ is required to satisfy the normalization condition

$$A = \sqrt{\frac{2}{\pi}} / \left(\sqrt{\frac{2}{\pi}} + \frac{\Gamma[-q + \frac{1}{2}]}{2 \times \frac{1}{2}(-q + \frac{1}{2})} \right). \quad (32)$$

Further explanations of the fitting procedure for these fitting parameters are described in Appendix A.

The mathematical form of Eq. (28) is similar to the ST mass function, which demands A, q to be constants. Here we extend the ST mass function to the conditional mass function and introduce $A(\kappa)$, $q(\kappa)$ and $\varepsilon(S_2, \kappa)$ as fitting functions. The definition of μ' here has been so motivated to fit the probability distribution derived from random walks.

However, it has been noted that when the time interval is small, i.e., z_2 close to z_1 , the concept of random walks fails to describe the merger events because the excursion set neglects the correlations between scales (Peacock & Heavens 1990, Sheth & Tormen 2001), and this explains why the conditional mass function derived from the moving-barrier in a recent report (Sheth & Tormen 2001) may not be applied to construct the merger tree. Although Eqs. (28) and (29) faithfully describe the result of our random-walk processes, in order to make the conditional mass function useful even at small time interval, we find that μ' should be modified into

$$\mu' \equiv \frac{[\delta_c(z_1) - \delta_c(z_2)]\varepsilon(S_2, \kappa)}{(S_1 - S_2 - \Delta\delta_c^2(z_2))^{1/2}}. \quad (33)$$

The replacement of $S_1(M_1) - S_2(M_2)$ by $S_1(M_1) - S_2(M_2) - \Delta\delta_c^2(z_2)$ means that there exists a small mass gap between parent halos and progenitor halos. This small mass gap only becomes important when the look-back time is small and becomes negligible for a large look-back time. A small positive Δ serves to correlate the probability of progenitor halos to that of parent halos. We find that $\Delta = 0.06$ gives the best result when compared with simulations, as discussed below. Furthermore, the variance and threshold value appear as $\sqrt{S}/\delta_c(z_2)$ in μ' , such that the active picture and the passive picture are consistent with each other.

4.2. Comparison with N-body simulation

In order to test whether the conditional mass function derived in the previous section is valid for describing the formation history of dark halos, we compare our conditional mass function with the results of the particle-particle/ particle-mesh simulation given by Somerville et al. (2000). This simulation is a run for the τ CDM model ($\Omega = 1, \Gamma = 0.21, h =$

0.5) with box size 170 Mpc, particle number 256³. The halos are identified by using the ‘friends-of-friends’ (FOF) algorithm.

To convert the conditional mass function into the probability that a halo with mass M_2 at redshift z_2 had a progenitor in the mass range M_1 to $M_1 + dM_1$ at redshift z_1 , we have used (LC93)

$$\frac{dn(M_1, z_1|M_2, z_2)}{dM_1}dM_1 = \frac{M_2}{M_1}f(S_1, \delta_c(z_1)|S_2, \delta_c(z_2))\left|\frac{dS_1}{dM_1}\right|dM_1, \quad (34)$$

where

$$f(S_1, \delta_c(z_1)|S_2, \delta_c(z_2)) = f(\mu')\frac{\partial\mu'}{\partial S_1}. \quad (35)$$

The mass range of the parent halo chosen by Somerville et al. is $7.9 \sim 12.6M_L$ at the present time ($z_2 = 0$), where $M_L \sim 5.0 \times 10^{11}M_\odot$. Because the parent mass range is wide, we adopt a fair assumption that the parents of different mass obey the ST mass function. We use this distribution of M_2 to weigh the conditional mass function for the parent population. In Fig. 4, we compare the conditional mass function predicted by our nonspherical collapse procedure and by EPS with the results of the τ CDM simulation mentioned above. The four panels correspond to 4 different redshifts: $z_1 = 0.2, 0.5, 1.1$ and 2.1 . The dashed curves denote the prediction of the weighed EPS theory, the solid curves show our prediction by using the suitable definition of μ' (Eq. (33)), and the results of τ CDM simulation are represented by filled triangles. The comparison shows that our conditional mass function (Eq. (28) & Eq. (33)) provides a better description than EPS does, no matter z_1 is small or large. The successful description of our conditional mass function indicates that the nonspherical collapse boundary is more “realistic” and it may be applied to construction of merger trees.

5. Conclusion

In this paper, we have attempted to show that the nonspherical collapse boundary (NCB) reproduces the mass function, bias factor and conditional mass function better than Press-Schechter theory. This model successfully addresses these subjects over a wide range of mass scales, while the (extend) Press-Schechter theory always fails in some mass range.

Through the simulations of 6-dimensional random walks, we have shown that the bias function generated by NCB agrees with the N-body simulations reasonably well. Because NCB is more complicated than the PS’s collapse boundary, the relation between the bias function and the conditional mass function is no longer as simple as the EPS model. Especially, the two issues are not of identical representation in the framework of NCB model. We provide a formula for the conditional mass function of the nonspherical collapse boundary,

which can be applied to any two redshifts. In order to overcome the problem that random-walks fail to describe the merger history at small look-back time, we add a minor correction in the definition of variable μ' (c.f. Eq. (33)) while maintaining the overall mathematical form intact. This formula can describe the conditional mass probability more accurately than the EPS formalism over a wide mass range, even when the two epochs are close. All the reasonable results we obtained from NCB suggest that NCB is sufficiently reasonable to reproduce the statistical properties of dark halos.

There are two interesting issues pertinent to NCB: one is its natural association with the halo angular momentum through its dependence on the quantity r . The halo angular momentum distribution predicted by the NCB model is reported elsewhere (Chiueh, Lee and Lin, in preparation). Second, as mentioned in § 4, the conditional mass function offers an useful tool to construct the merger history through merger trees. Recent works have indicated that the merger history constructed by using EPS's conditional mass function deviates from that extracted from N-body simulations, and it is thought to be caused mainly by the spherical-symmetry assumption of the EPS model rather than the specific features of the merger tree scheme. Thus, we expect the disagreement can be greatly improved by the NCB conditional mass function given in this work.

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A. The fitting procedure of conditional mass function

After performing the random walks between two collapse boundaries, which are mentioned in § 4.1, we use the mathematical form of ST mass function and tune the values of A , q and ε to adjust our results of random walks. It should be emphasized that A and q are not tuned independently because the integral of the function $f(\mu')$ should be normalized to unity, due to the assumption that all the trajectories starting from the first boundary will eventually hit the second boundary. Thus, we have

$$\int f(\mu')d\mu' = \int 2A(\kappa)(1 + \frac{1}{\mu'^{2q}})(\frac{1}{2\pi})^{1/2}\exp(-\frac{\mu'^2}{2})d\mu' = 1, \quad (A1)$$

then we get

$$A = \sqrt{\frac{2}{\pi}} / (\sqrt{\frac{2}{\pi}} + \frac{\Gamma[-q + \frac{1}{2}]}{2 \times \frac{1}{2}(-q + \frac{1}{2})}). \quad (A2)$$

In addition, it is expected that all three parameters A , q and ε should yield a fitting formula of conditional probability to match the two extreme cases, as mentioned in § 4.1. For κ approaching unit, i.e., the two boundaries are very close, $A(\kappa)$ gives the value $\frac{1}{2}$, and q is thus zero, which agrees with the Press-Schechter’s conditional probability. For κ to be very large, A and q approach 0.322 and 0.3 respectively, the values given by Sheth & Tormen (1999). The results of our 6-dimensional simulation and our fitting formula Eq. (28) and (29) are shown in Fig. 5 ~ Fig. 9. The five figures correspond to different separations of boundaries, i.e., $\kappa = 1.1, 1.3, 2.0, 3.0$, and 4.0 . The top panel of each figure is the result for initial mass scale $\sqrt{S_2} = 0.7$, the middle panel is for $\sqrt{S_2} = 1.5$ (around M^*) and the bottom is for $\sqrt{S_2} = 2.9$. The opaque circles represent the data from 6-dimensional simulation and the solid lines are the best fit using Eq. (28).

We find that for the case κ is fixed, A does not vary with the initial mass scale S_2 , i.e., A is a function of κ only. In Fig. 10, we show the relation between A and $1/\kappa$. The square symbols show the fitting values of A for five different κ plus two boundary values: $(1/\kappa, A) = (0, 0.322)$ and $(1, 0.5)$, which recover the limit conditions associated with EPS and ST cases. The solid line represents here is a linear relation between A and $1/\kappa$:

$$A = 0.322 + \frac{0.178}{\kappa}, \quad (\text{A3})$$

which is determined by the above two boundary points. This line almost passes through the other five fitting points and thus is a good representation for $A(\kappa)$. Combined with Eq. (A2), q can be determined immediately for each κ .

Our tests also show that ε is a function of a combined variable $x \equiv (\sqrt{S_2}/\delta_c(z_2) - 0.25)/\kappa$. The best fit of the function is (see Fig. 11)

$$\varepsilon(x) = 0.036x^4 - 0.309x^3 + 0.944x^2 - 1.060x + 1. \quad (\text{A4})$$

This result is consistent with the geometry of two boundaries. Since we have the r -distribution of various mass scales S_2 ’s that hit the first boundary, one can estimate roughly the distance between two boundaries for each mass scale. It is easily seen that for a fixed κ , the gap d initially decreases to reach a minimum and then rises when S_2 continues to increase. The behavior of $d/(\delta_c(z_1) - \delta_c(z_2))$ agrees with that of ε defined above.

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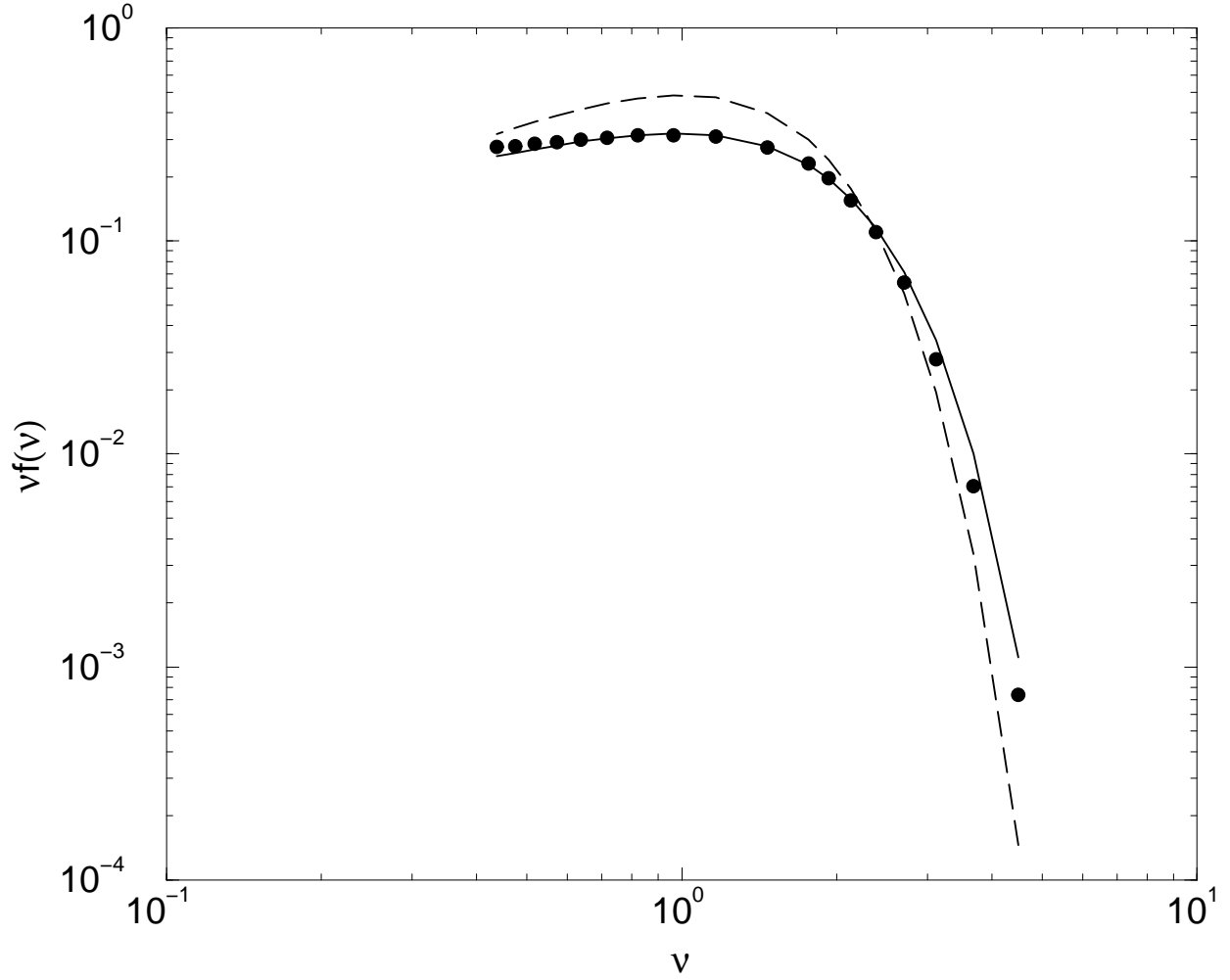


Fig. 1.— Comparison of the mass function derived from EPS model (dashed line) and the nonspherical collapse boundary (filled circles) with Sheth & Tormen’s fitting formula of N-body simulations (solid line).

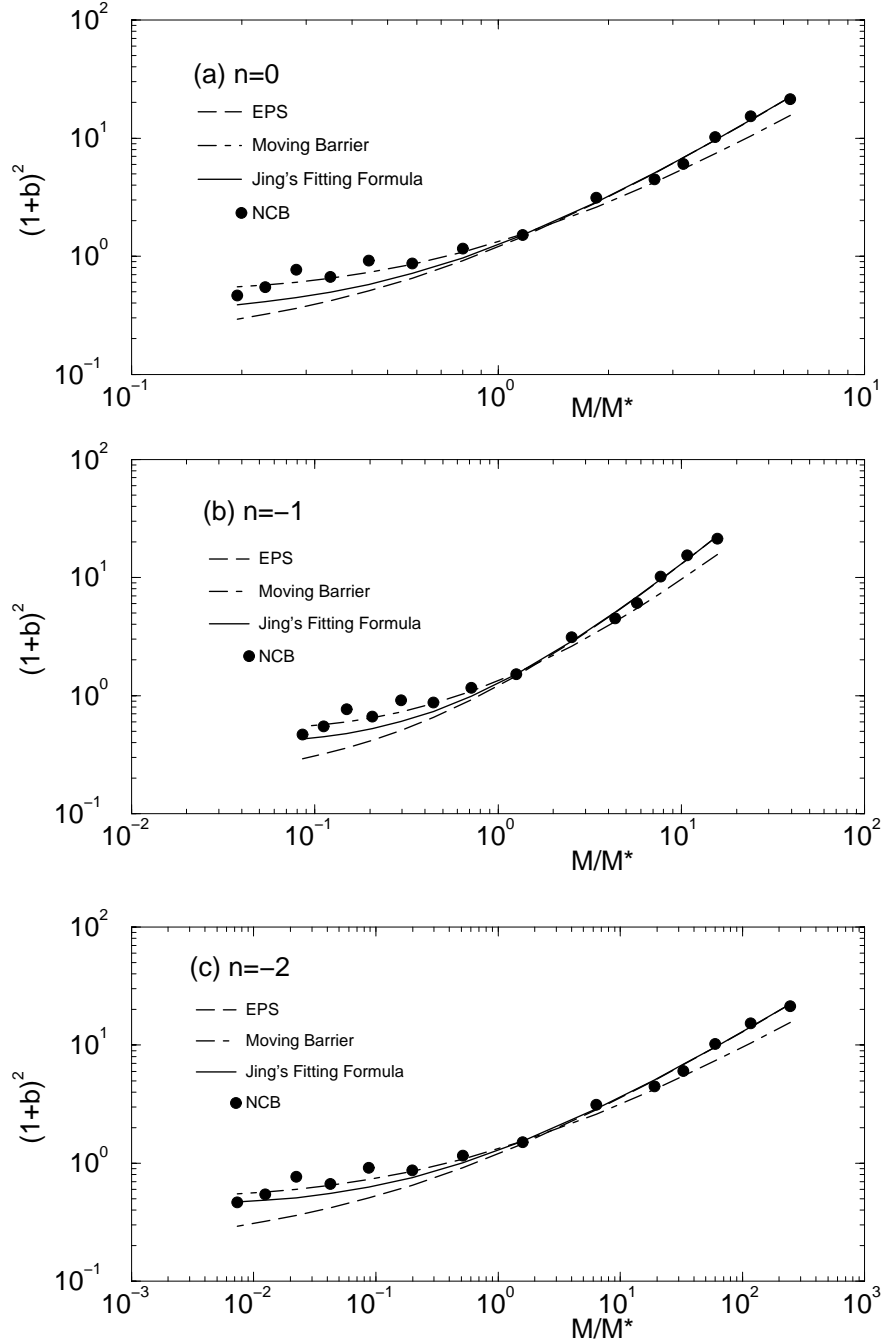


Fig. 2.— The bias factor as a function of M/M^* for three different initial power indices: (a) $n=0$, (b) $n=-1$ and (c) $n=-2$. The circle symbols show the results of our algorithm of Y -space random walks. The solid curves are Jing's fitting formula derived from N -body simulations. The long dashed lines denote the analytical bias prediction based on EPS (MW96), and the dot-dashed lines show the results predicted by using moving-barrier model (Sheth, Mo & Tormen 2001).

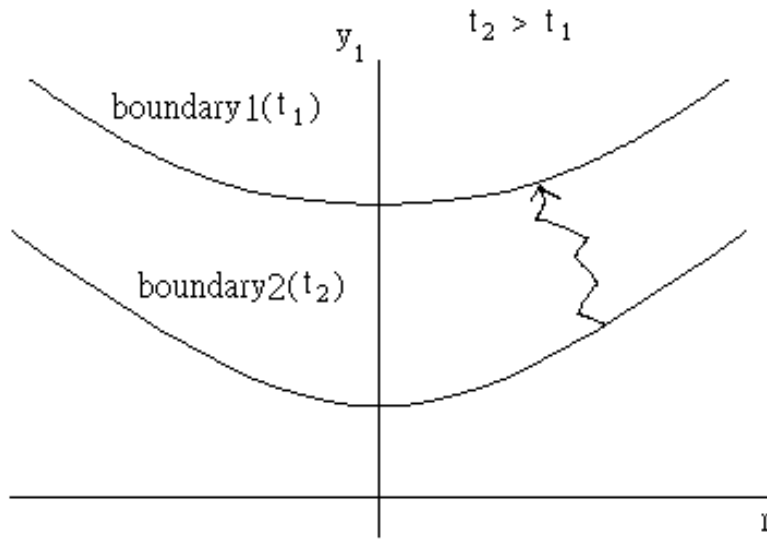


Fig. 3.— The nonspherical collapse boundaries of two different epochs. The higher the boundary is, the earlier the collapse time.

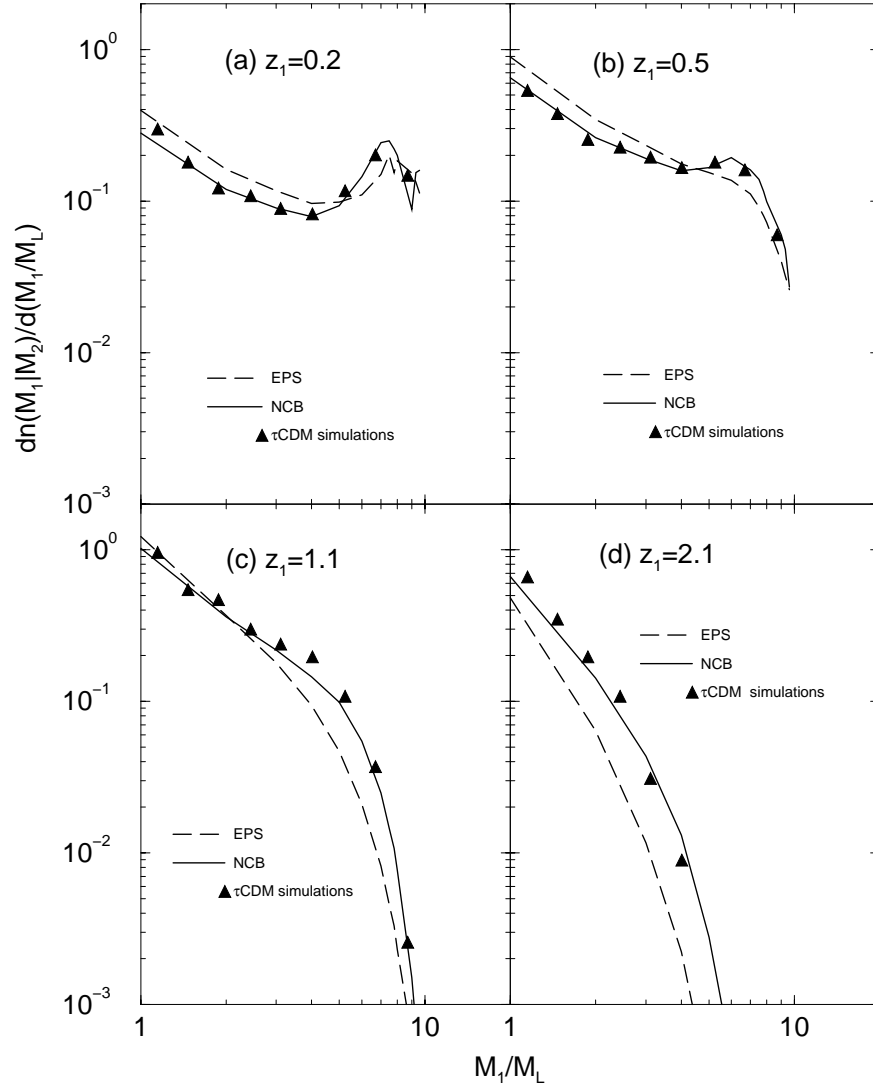


Fig. 4.— Comparison of the conditional mass function predicted by EPS (dashed curves) and NCB (solid curves) models with the results of the τ CDM simulations (triangle symbols). The parent halos are chosen at $z = 0$, and the progenitors are at (a) $z = 0.2$, (b) $z = 0.5$, (c) $z = 1.1$ and (d) $z = 2.1$.

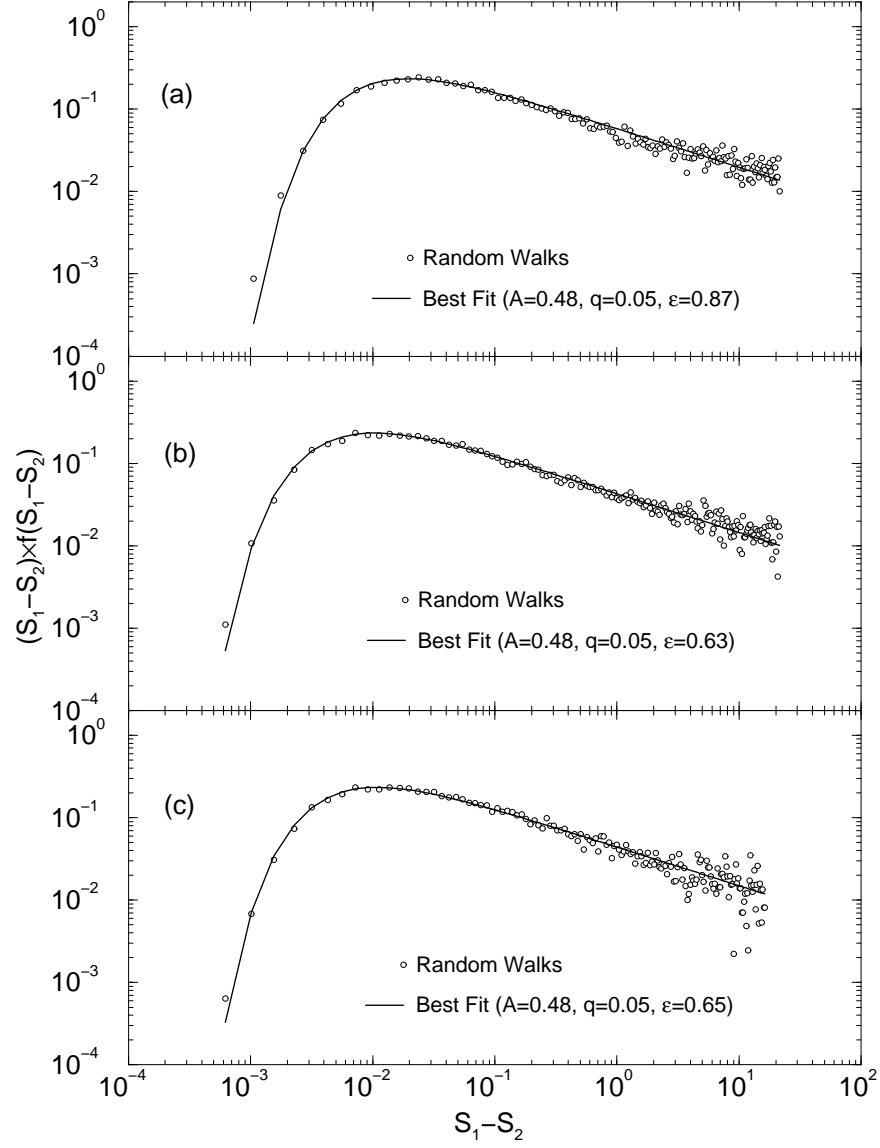


Fig. 5.— The conditional mass function for $\kappa = 1.1$. The opaque circles represent the results of the six-dimensional random walks and the solid lines are the best fit using Eqs. (28) and (29). The mass scale at time z_2 is chosen as (a) $\sqrt{S_2} = 0.7$, (b) $\sqrt{S_2} = 1.5$ and (c) $\sqrt{S_2} = 2.9$.

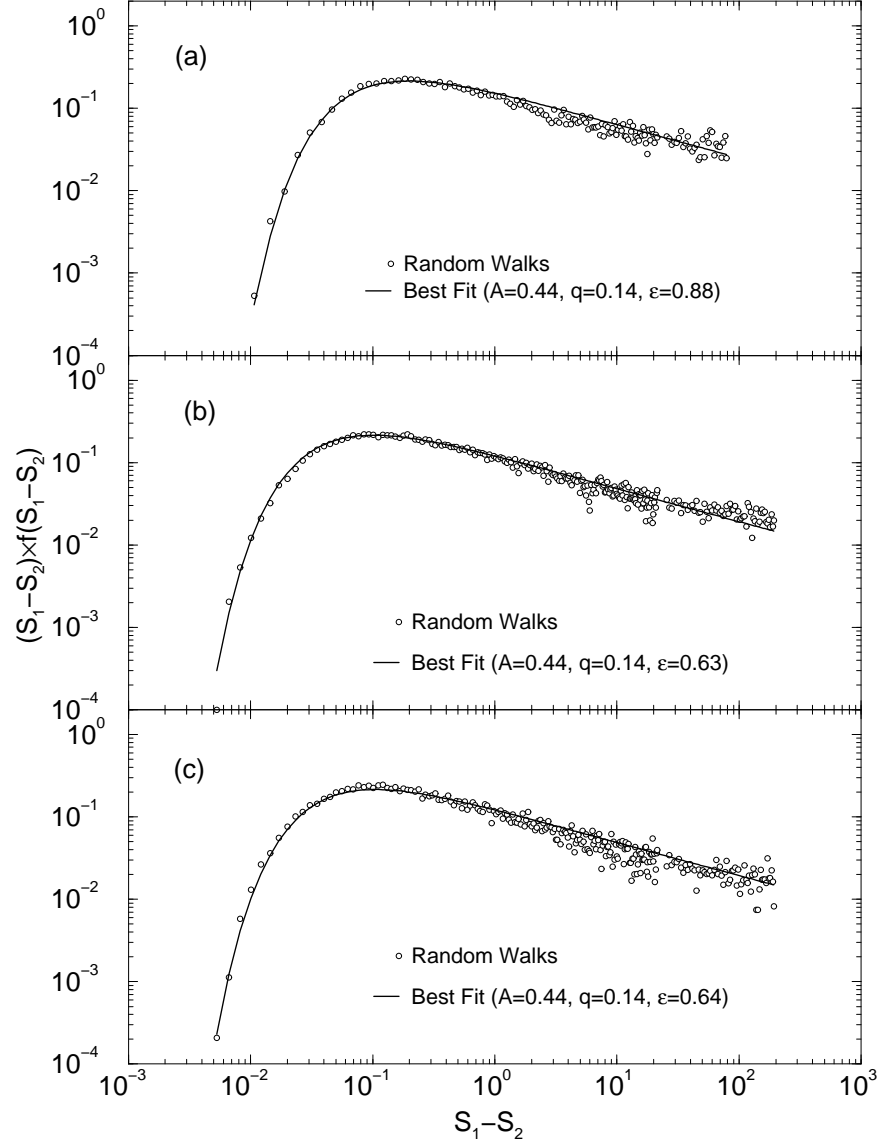


Fig. 6.— Same as the previous figure, but for $\kappa = 1.3$.

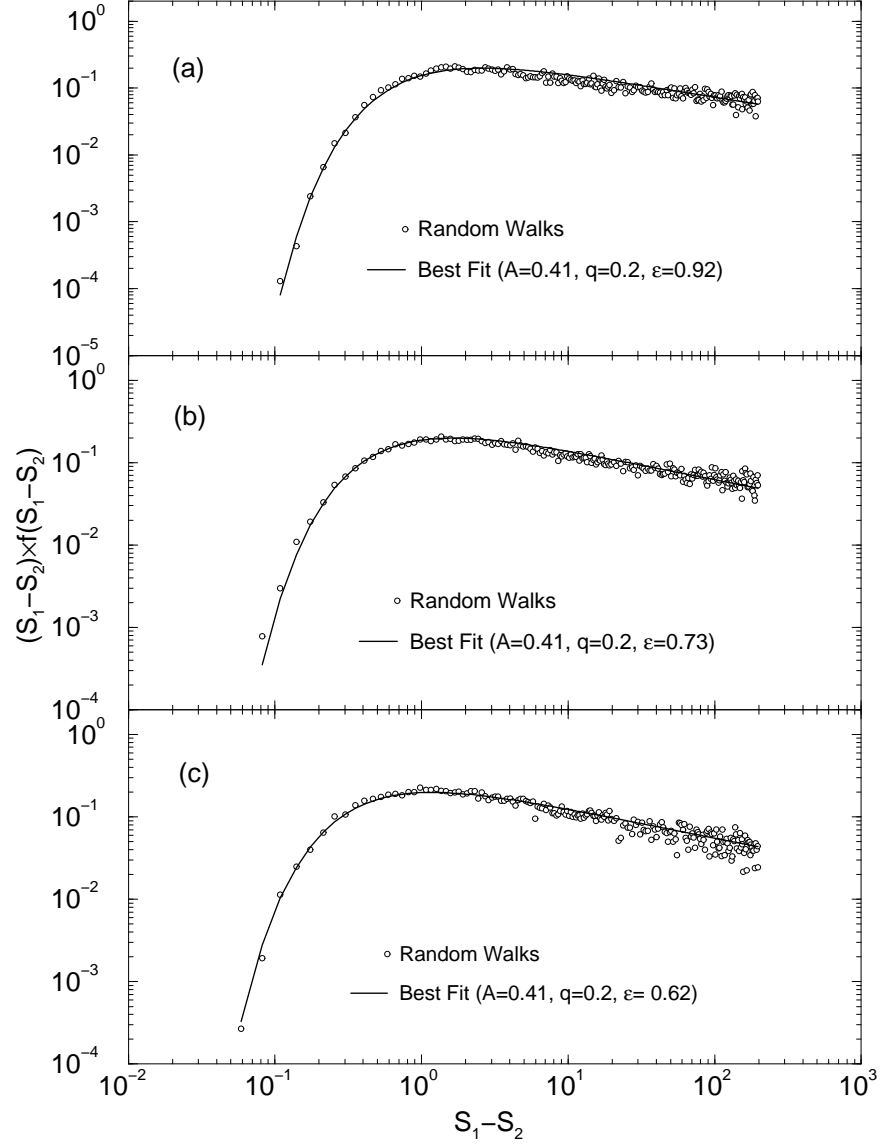


Fig. 7.— Same as the previous figure, but for $\kappa = 2.0$.

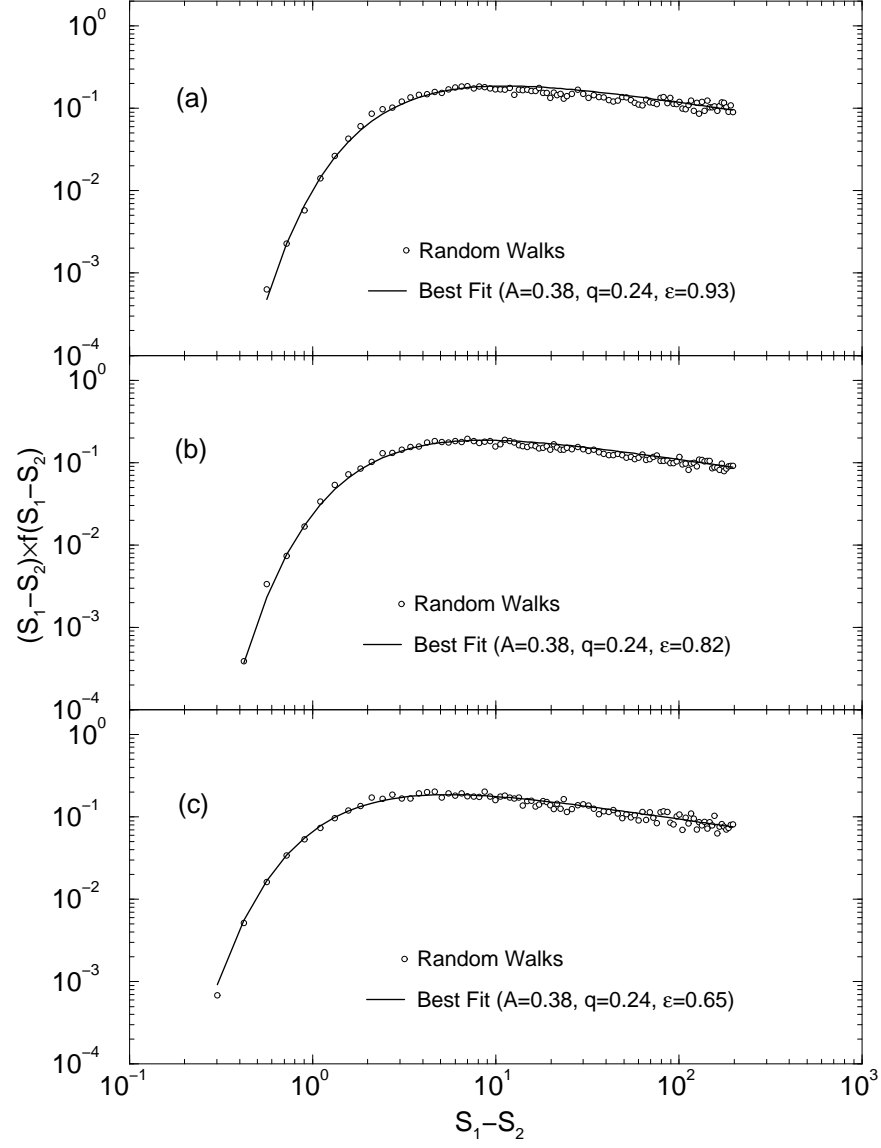


Fig. 8.— Same as the previous figure, but for $\kappa = 3.0$.

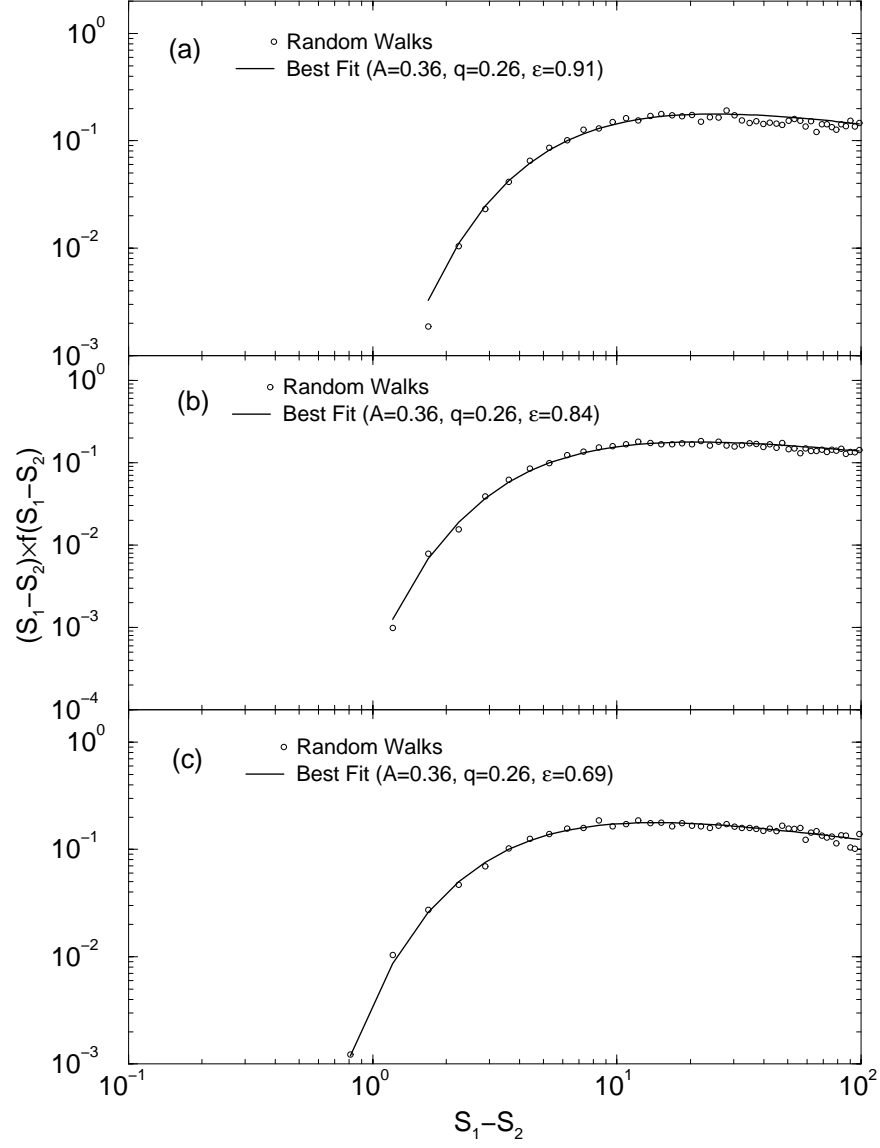


Fig. 9.— Same as the previous figure, but for $\kappa = 4.0$.

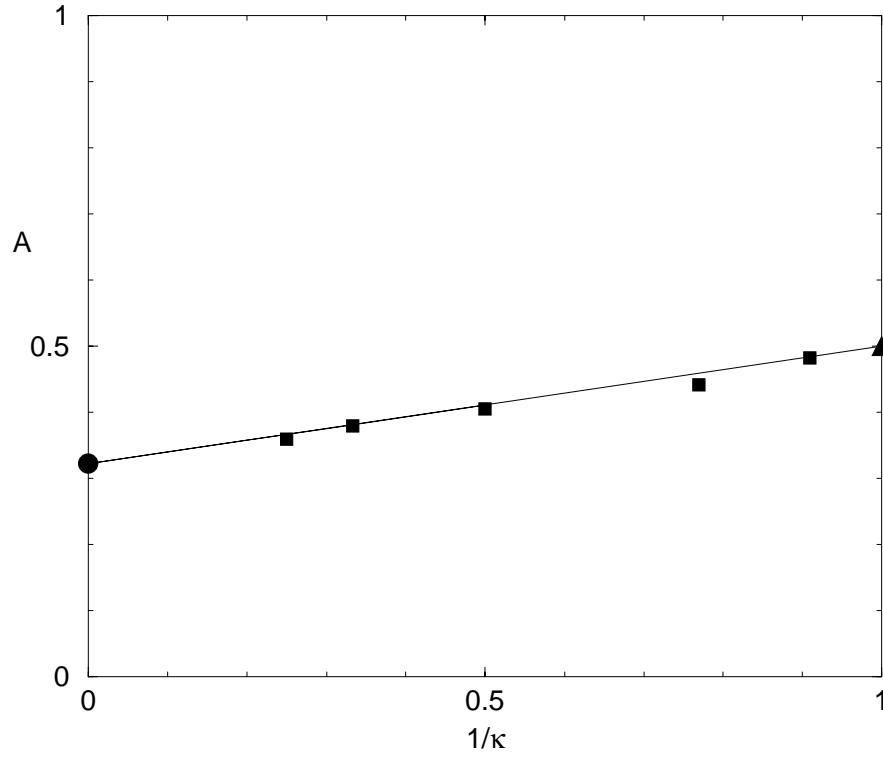


Fig. 10.— $A - 1/\kappa$ relation. The square symbols show the five fitting values of A for our data, the filled triangle represents the EPS case, the filled circle represents the case of ST mass function, and the solid line denotes the linear relation between A and $1/\kappa$. (Eq. (30))

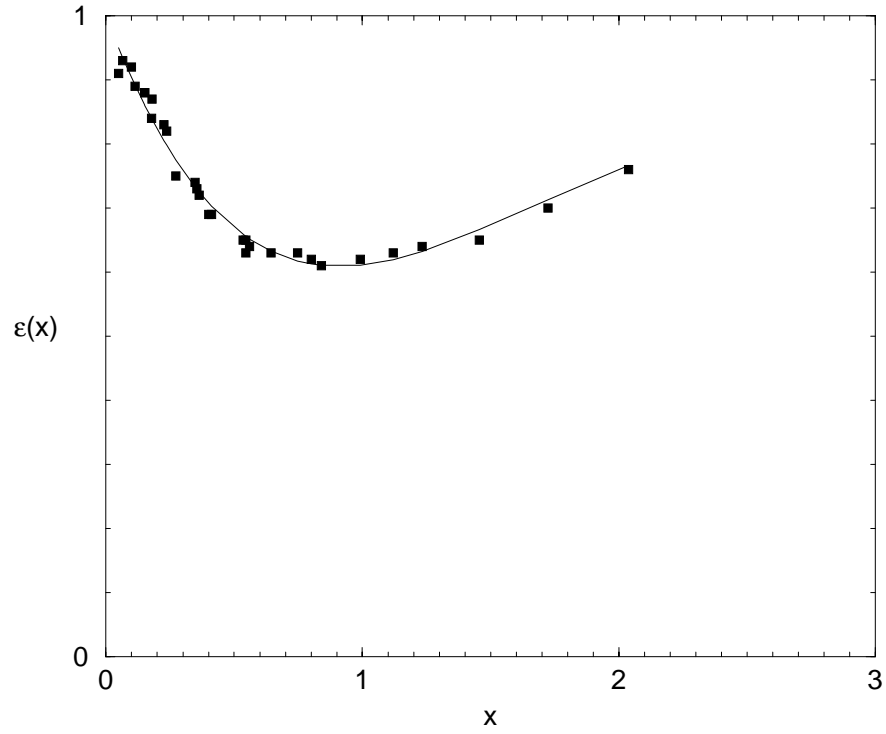


Fig. 11.— ε factor as a function of $x(\equiv (\sqrt{S_2}/\delta_c(z_2) - 0.25)/\kappa)$. Similar to Fig. 10, the square symbols show the fitting values of ε for our data and the solid line denotes the fitting relation of the square symbols. (Eq. (31))